

TWINS Equatorial Pitch Angle Averaged Ion Flux

The deconvolved 40 keV equatorial pitch angle averaged ion flux in units of $[\text{eV cm}^2 \text{sr s}]^{-1}$. The sun is to the left. The radius extends to $8 R_E$. Distances are marked in $1 R_E$ intervals by gray dashed circles. The principal peak of the ion flux is marked by a black star.

The technique used to extract the equatorial ion pitch angle distributions from the ENA images is described in Appendix A of *Perez, et al.* [2012]. In this method, the ion equatorial pitch angle distribution is expanded in a linear combination of tri-cubic splines [*deBoor, 1978*]. The expansion coefficients are then obtained by minimizing a combination of normalized chi-squared and a penalty function derived by *Wahba* [1990]. Requiring that normalized chi-squared is near unity ensures that the resulting distribution fits the data. Including the penalty function in the minimization ensures that the result is as smooth (in the sense of a minimum second derivative) as is consistent with fitting the data. In this procedure, spatial structure is minimized and appears in the result only to the extent that it is necessary. Thus, while there may be more and smaller scale structure that is not resolved, the structure that is found is statistically required to fit the data, i.e., match the ENA images.

The uncertainties in each pixel of the ENA image are a statistical measure of the information content of the data. The second moment of the 15-16 individual sweeps is used to estimate the uncertainties in each pixel of the time-integrated image.

In order to deconvolve the ion distributions, magnetic field mapping is required. For this study the *Tsyganenko and Sitnov* [2005] magnetic field model was used. The density of neutral hydrogen, i.e., the geocorona, is also needed. The TWINS exospheric neutral hydrogen density model was used [*Zoennchen, et al. 2013*]. To include the LAEs (Low Altitude Emissions), the thick target approximation of *Bazell et al.* [2010] was used.

References:

- Bazell, D., E. C. Roelof, T. Sotirelis, P. C. Brandt, H. Nair, P. Valek, J. Goldstein, and D. McComas (2010), Comparison of TWINS images of low-altitude emission of energetic neutral atoms with DMSP precipitating ion fluxes, *J. Geophys. Res.*, *115*, A10204, doi:10.1029/2010JA015644.
- deBoor, C. (1978), *A Practical Guide to Splines*, Springer-Verlag, N.Y.
- Perez, J. D., E. W. Grimes, J. Goldstein, D. J. McComas, P. Valek, and N. Billor (2012), Evolution of CIR storm on 22 July 2009, *J. Geophys. Res.*, *117*, A09221, doi:10.1029/2012JA017572.
- Tsyganenko, N. A., and M. I. Sitnov (2005), Modeling the dynamics of the inner magnetosphere during strong geomagnetic storms, *J. Geophys. Res.*, *110*, A03208, doi:10.1029/2004JA010798.
- Wahba, G. (1990), *Spline Models for Observational Data*, Soc. for Ind. and Appl. Math., Philadelphia, Pa., doi:10.1137/1.9781611970128.
- Zoennchen, J. H. and U. Nass and H. J. Fahr (2013), Exospheric hydrogen density distributions for equinox and summer solstice observed with TWINS1/2 during solar minimum, *Ann. Geophys.*, *31*, 513, doi:10.5194/angeo-31-513-2013.